

AN EXCEPTIONAL VHE GAMMA-RAY FLARE OF PKS2155–304

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ABSTRACT

The high-frequency peaked BL Lac PKS 2155–304 at redshift $z = 0.116$ is a well-known VHE (>100 GeV) γ -ray emitter. Since 2002 its VHE flux has been monitored using the H.E.S.S. stereoscopic array of imaging atmospheric-Cherenkov telescopes in Namibia. During the July 2006 dark period, the average VHE flux was measured to be more than ten times typical values observed from the object. This article focuses solely on an extreme γ -ray outburst detected in the early hours of July 28, 2006 (MJD 53944). The average flux observed during this outburst is $I(>200 \text{ GeV}) = (1.72 \pm 0.05_{\text{stat}} \pm 0.34_{\text{syst}}) \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$, corresponding to ~ 7 times the flux, $I(>200 \text{ GeV})$, observed from the Crab Nebula. Peak fluxes are measured with one-minute time scale resolution at more than twice this average value. Variability is seen up to ~ 600 s in the Fourier power spectrum, and well-resolved bursts varying on time scales of ~ 200 seconds are observed. There are no strong indications for spectral variability within the data. Assuming the emission region has a size comparable to the Schwarzschild radius of a $\sim 10^9 M_{\odot}$ black hole, Doppler factors greater than 100 are required to accommodate the observed variability time scales.

Subject headings: Galaxies: active – BL Lacertae objects: Individual: PKS 2155–304 – Gamma rays: observations

1. INTRODUCTION

Flux variability studies provide a strong probe into the physical processes of the innermost regions of Ac-

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tive Galactic Nuclei (AGN). Although the broad-band emission from all AGN is highly variable, the most extreme flux variability, i.e. largest magnitude and shortest time scale, is observed from a class of AGN known as blazars. As a result blazar variability studies are crucial to unraveling the mysteries of AGN. Over a dozen blazars have been detected so far at VHE energies. In the Southern Hemisphere, PKS 2155–304 is generally the brightest blazar at these energies, and is probably the best-studied at all wavelengths. The VHE flux observed (Aharonian et al. 2005a) from PKS 2155–304 is typically of the order $\sim 15\%$ of the Crab Nebula flux above 200 GeV. The highest flux previously measured in one night is approximately four times this value and clear VHE-flux variability has been observed on daily time scales. The most rapid flux variability measured for this source is 25 min (Aharonian et al. 2005b), occurring at X-ray energies. The fastest variation published from any blazar, at any wavelength, is an event lasting ~ 800 s where the X-ray flux from Mkn 501 varied by 30% (Xue & Cui 2005)³⁰, while at VHE energies doubling timescales as fast as ~ 15 minutes have been observed from Mkn 421 (Gaidos et al. 1996).

The High Energy Stereoscopic System (H.E.S.S.; Hinton 2004) is used to study VHE γ -ray emission from wide variety of astrophysical objects. As part of the normal H.E.S.S. observation program the flux from known VHE AGN is monitored regularly to search for bright flares. During such flares, the unprecedented sensitivity of H.E.S.S. (5 standard deviation, σ , detection in ~ 30 s for a Crab Nebula flux source at 20° zenith angle) enables studies of VHE-flux variability on time scales of a few tens of seconds. During the July 2006 dark period, the average VHE flux observed by H.E.S.S. from PKS 2155–304 was more than ten times its typical value. In particular, an extremely bright flare of PKS 2155–304 was observed in the early hours of July 28, 2006 (MJD 53944). This article focuses solely on this particular flare. The results from other H.E.S.S. observations of PKS 2155–304 from 2004 through 2006 will be published elsewhere.

2. RESULTS FROM MJD 53944

A total of three observation runs (~ 28 min each) were taken on PKS 2155–304 in the early hours³¹ of MJD 53944. These data entirely pass the standard H.E.S.S. data-quality selection criteria, yielding an exposure of 1.32 h live time at a mean zenith angle of 13° . The standard H.E.S.S. calibration (Aharonian et al. 2004) and analysis tools (Benbow 2005) are used to extract the results shown here. As the observed signal is exceptionally strong, the event-selection criteria (Benbow 2005) are performed using the *loose cuts*, instead of the *standard cuts*, yielding an average post-analysis energy threshold of 170 GeV. The *loose cuts* are selected since they have a lower energy threshold and higher γ -ray and background acceptance. The higher acceptances avoid low-statistics issues with estimating the background and significance on short time scales, thus simplifying the analysis. The

on-source data are taken from a circular region of radius $\theta_{cut} = 0.2^\circ$ centered on PKS 2155–304, and the background (off-source data) is estimated using the *Reflected-Region* method (Berge et al. 2007).

A total of 12480 on-source events and 3296 off-source events are measured with an on-off normalization of 0.215. The observed excess is 11771 events (~ 2.5 Hz), corresponding to a significance of 168σ calculated following the method of Equation (17) in Li & Ma (1983). It should be noted that use of the *standard cuts* also yields a strong excess (6040 events, 159σ) and results (i.e. flux, spectrum, variability) consistent with those detailed later.

2.1. Flux Variability

The average integral flux above 200 GeV observed from PKS 2155–304 is $I(>200 \text{ GeV}) = (1.72 \pm 0.05_{\text{stat}} \pm 0.34_{\text{syst}}) \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$, equivalent to ~ 7 times the $I(>200 \text{ GeV})$ observed from the Crab Nebula (I_{Crab} ; Aharonian et al. 2006). Figure 1 shows $I(>200 \text{ GeV})$, binned in one-minute intervals, versus time. The fluxes in this light curve range from $0.65 I_{\text{Crab}}$ to $15.1 I_{\text{Crab}}$, and their fractional root mean square (rms) variability amplitude (Vaughan et al. 2003) is $F_{\text{var}} = 0.58 \pm 0.03$. This is ~ 2 times higher than archival X-ray variability (Zhang et al. 1999; Zhang et al. 2005). The Fourier power spectrum calculated from Figure 1 is shown in Figure 2. The error on the power spectrum is the 90% confidence interval estimated from 10^4 simulated light curves. These curves are generated by adding a random constant to each individual flux point, where this constant is taken randomly from a Gaussian distribution with a dispersion equal to the error of the respective point. The average power expected when the measurement error dominates is shown as a dashed line (see the Appendix in Vaughan et al. 2003). There is power significantly above the measurement noise level up to $1.6 \times 10^{-3} \text{ Hz}$ (600 s). The power spectrum also shows that most of the power is at low frequencies. The grey shaded area shows the 90% confidence level obtained by simulating 10^4 light curves with a power-law Fourier spectrum $P_\nu \propto \nu^{-2}$ (Timmer & Koenig 1995) and a random Gaussian error as above. The power spectrum derived from the data is thus compatible with a light curve generated by a stochastic process with a power-law Fourier spectrum of index -2. An index of -1 produces too much power at high frequencies and is rejected. These power spectra are remarkably similar to those derived in X-rays (Zhang et al. 1999) from the same source.

Rapid variability is clearly visible in substructures that appear in the light curve, with even shorter rise and decay time scales than found in the Fourier analysis. In order to quantify those time scales, the light curve is considered as consisting of a series of bursts, which is common for AGN and γ -ray bursts (GRBs). The “generalized Gaussian” shape from Norris et al. (1996) is used to characterize these bursts, where the burst intensity is described by: $I(t) = A \exp[-(|t - t_{\text{max}}|/\sigma_{r,d})^\kappa]$, where t_{max} is the time of the burst’s maximum intensity (A); σ_r and σ_d are the rise ($t < t_{\text{max}}$) and decay ($t > t_{\text{max}}$) time constants, respectively; and κ is a measure of the burst’s sharpness. The rise and decay times, from half to maximum amplitude, are $\tau_{r,d} = [\ln 2]^{1/\kappa} \sigma_{r,d}$. A peak finding

³⁰ Xue & Cui (2005) also demonstrate that a 60% X-ray flux increase in ~ 200 s observed (Catanese & Sambruna 2000) from Mkn 501 is likely an artifact.

³¹ The three runs began at 00:35, 01:06 and 01:36 UTC, respectively.

TABLE 1

THE RESULTS OF THE BEST χ^2 FIT OF THE SUPERPOSITION OF FIVE BURSTS AND A CONSTANT TO THE DATA SHOWN IN FIGURE 1. THE CONSTANT TERM IS $0.27 \pm 0.03 \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$ ($1.1 I_{\text{Crab}}$).

t_{max} [min]	A [$10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$]	τ_r [s]	τ_d [s]	κ
41.0	2.7 ± 0.2	173 ± 28	610 ± 129	1.07 ± 0.20
58.8	2.1 ± 0.9	116 ± 53	178 ± 146	1.43 ± 0.83
71.3	3.1 ± 0.3	404 ± 219	269 ± 158	1.59 ± 0.42
79.5	2.0 ± 0.8	178 ± 55	657 ± 268	2.01 ± 0.87
88.3	1.5 ± 0.5	67 ± 44	620 ± 75	2.44 ± 0.41

tool, using a Markov chain algorithm (Morhac 2000), selected five significant bursts. A function consisting of a superposition of an identical number of bursts plus a constant signal was fit³² to the data. The best fit has a χ^2 probability of 20% and the fit parameters are shown in Table 1. Interestingly, there is a marginal trend for κ to increase with subsequent bursts, making them less sharp, as the flare progresses. The κ values are close to the bulk of those found by Norris et al. (1996), but the time scales measured here are two orders of magnitude larger.

During both the first two bursts there is clear doubling of the flux within τ_r . Such doubling is sometimes used as a characteristic time scale of flux variability. For compatibility with such estimators, the definition of doubling time, $T_2 = |I_{ij}\Delta T/\Delta I|$, from Zhang et al. (1999) was used³³. Here, $\Delta T = T_j - T_i$, $\Delta I = I_j - I_i$, $I_{ij} = (I_j + I_i)/2$, with T and I being the time and flux, respectively, of any pair of points in the light curve. The fastest $T_2 = 224 \pm 60$ s is compatible with the fastest significant time scale found by the Fourier transform. Averaging the five lowest T_2 values yields 330 ± 40 s.

The variability time scales of these bursts are among (see also Albert et al. 2007) the fastest ever seen in a blazar, at any wavelength, and are almost an order of magnitude smaller than previously observed from this object. It should be noted that similar time scales are found with even smaller binning (e.g. 20 s) of the H.E.S.S. light curve, and that many checks of the data quality were undertaken to ensure that the flux variations cannot be the result of background fluctuations, atmospheric events, etc. In addition, all the results have been verified using an independent calibration method and alternative analysis techniques.

2.2. Spectral Analysis

Figure 3 shows the time-averaged photon spectrum for these data. The data are well fit, $\chi^2 = 17.1$ for 13 degrees of freedom (d.o.f.), by a broken power-law function:

$$E < E_B : \frac{dN}{dE} = I_0 \left(\frac{E}{1 \text{ TeV}} \right)^{-\Gamma_1}$$

$$E > E_B : \frac{dN}{dE} = I_0 \left(\frac{E_B}{1 \text{ TeV}} \right)^{(\Gamma_2 - \Gamma_1)} \left(\frac{E}{1 \text{ TeV}} \right)^{-\Gamma_2},$$

where $I_0 = (2.06 \pm 0.16 \pm 0.41) \times 10^{-10} \text{ cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$, $E_B = 430 \pm 22 \pm 80 \text{ GeV}$, $\Gamma_1 = 2.71 \pm 0.06 \pm 0.10$, and $\Gamma_2 = 3.53 \pm 0.05 \pm 0.10$. For each parameter, the two uncertainties are the statistical and systematic values, respectively. Fits to the data of either a simple

power law ($\Gamma = 3.19 \pm 0.02 \pm 0.10$, $\chi^2 = 138$, 15 d.o.f) or a power law with an exponential cut-off ($\chi^2 = 45$, 14 d.o.f.) are not acceptable. The time-averaged spectrum ($\Gamma = 3.32$) of PKS 2155–304 measured in 2003 (Aharonian et al. 2005a), multiplied by the ratio (48.7) of $I(>200 \text{ GeV})$ from the respective data sets, is also shown in Figure 3. Despite a factor of ~ 50 change in flux there is qualitatively little difference between the two spectra. Indeed, fitting a broken power law to the current data set, keeping Γ_1 and Γ_2 fixed to the values measured in 2003, yields a value for E_B consistent with that measured in 2003. The small difference is surprising since a change of the spectral shape with varying flux levels, typically hardening with increased flux, has often been observed from blazars at X-ray energies (see, e.g., Giommi et al. 1990), as well as in the VHE domain (see, e.g., Aharonian et al. 2002).

The high flux observed from PKS 2155–304 allows the determination of accurate photon spectra on time scales of the order of minutes. Therefore, a simple search for temporal changes of the VHE spectral shape within these data was performed. Spectra were determined for consecutive data slices of 28 minutes (1 run), 10 minutes, and 5 minutes. Fitting the time-average spectral shape, allowing only the normalization (I_0) to vary, to these short-time-scale spectra yields reasonable χ^2 probabilities. Thus, there are no strong indications of fast spectral variability. However, weak variations ($\Delta\Gamma < 0.2$) are not ruled out. A more sophisticated study of any fast spectral variations within these data is beyond the scope of this letter and will be published elsewhere.

3. DISCUSSION

It is very likely that the electromagnetic emission in blazars is generated in jets that are beamed and Doppler-boosted toward the observer. Superluminal expansions observed with VLBI (Piner & Edwards 2004) provide evidence for moderate Doppler boosting in PKS 2155–304. Causality implies that γ -ray variability on a time scale t_{var} , with a Doppler factor³⁴ (δ), is related to the radius (R) of the emission zone by $R \leq ct_{\text{var}}\delta/(1+z)$. Conservatively using the best-determined rise time (i.e. τ_r with the smallest error) from Table 1 for $t_{\text{var}} = 173 \pm 28$ s (note that this is similar to the fastest T_2) limits the size of the emission region to $R\delta^{-1} \leq 4.65 \times 10^{12} \text{ cm} \leq 0.31 \text{ AU}$.

The jets of blazars are believed to be powered by accretion onto a supermassive black hole (SMBH). Thus accretion/ejection properties are usually presumed to scale with the Schwarzschild radius R_S of the SMBH, where $R_S = 2GM/c^2$, which is the smallest, most-natural size of the system (see, e.g., Blandford et al. 1982). Expressing the size R of the γ -ray emitting region in terms of R_S , the variability time scale limits its mass by $M \leq (c^3 t_{\text{var}} \delta / 2G(1+z)) R_S / R \sim 1.6 \times 10^7 M_\odot \delta R_S / R$. The reported³⁵ host galaxy luminosity $M_R = -24.4$ (Table 3 in Kotilainen et al. 1998) would imply a SMBH mass of order $1-2 \times 10^9 M_\odot$ (Bettoni et al. 2003), and therefore, $\delta \geq 60 - 120 R/R_S$. Emission regions of only a few R_S

³² The Markov chain burst positions were used to initialize t_{max} for each burst. All parameters are left free in the fit.

³³ Only values of T_2 with less than 30% uncertainty are considered.

³⁴ With δ defined in the standard way as $[\Gamma(1-\beta \cos \theta)]^{-1}$, where Γ is the bulk Lorentz factor of the plasma in the jet, $\beta = v/c$, and θ is the angle to the line of sight.

³⁵ See Wurtz et al. (1996) and $M_R > -23.1$ (for $h = 0.5$) showing the need for confirmation of this value.

would require values of δ much greater than those typically derived for blazars ($\delta \sim 10$) and come close to those used for GRBs, which would be a challenge to understand. For example, the sub-parsec VHE γ -ray emitting plasma would have to decelerate with a high efficiency to accommodate relatively small Lorentz factors observed at parsec scales (Piner & Edwards 2004). It is however possible that the SMBH mass is over-estimated, reducing the δ constraint by the same factor, or that the variability has an origin (e.g., a geometric effect from jet bending as discussed in Wagner et al. 1993) unrelated to the black hole. Detailed modeling of the spectral energy distribution of PKS 2155–304, during the multiple VHE flares observed by H.E.S.S. in the July 2006 dark period, including simultaneous multi-frequency data, will appear elsewhere.

The VHE variability observed in this particular flaring episode is the fastest ever observed from a blazar. While the variability is a factor of five times faster than previously measured from Mkn 421 (Gaidos et al. 1996), in terms of the light-crossing time of the Schwarzschild radius, R_S/c , the variability of PKS 2155–304 is another factor of ≈ 6 –12 more constraining assuming a $10^{8.22} M_\odot$ for Mkn 421 (Woo et al. 2005). It should also be noted that the choice of a ~ 3 minute variability time scale here is conservative and that the light curve is strongly over-sampled, allowing for the first time in the VHE regime

a detailed statistical analysis of a flare, which shows remarkable similarity to other longer duration events at X-ray energies. From such rapid variability one must conclude that either very large Doppler factors can be present in AGN jets, or that the observed variability is not connected to the central black hole, clearly showing the power of Cherenkov-telescope arrays in probing the internal mechanisms in BL Lacs.

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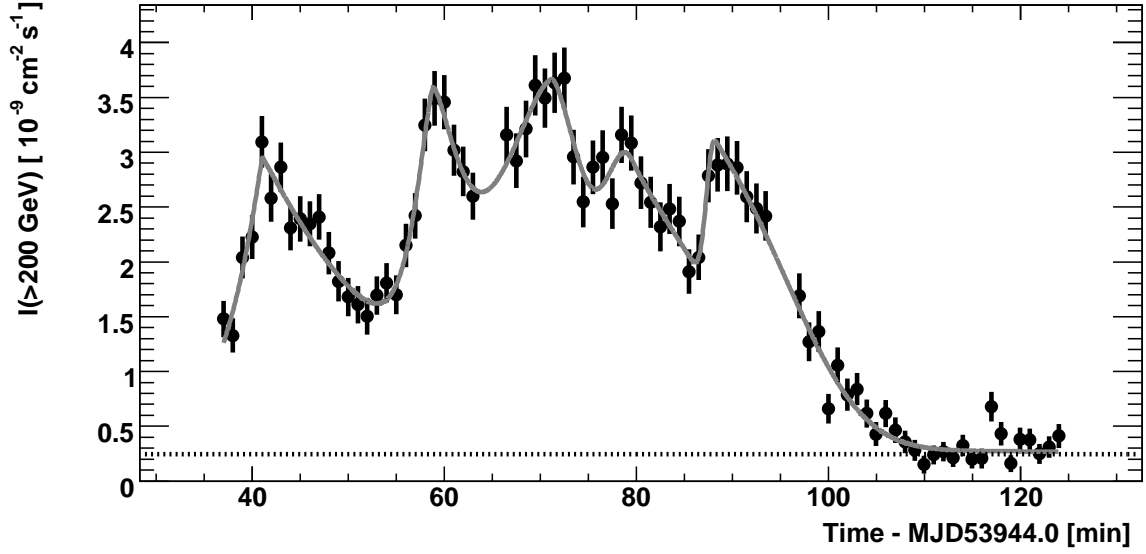


FIG. 1.— The integral flux above 200 GeV observed from PKS 2155–304 on MJD 53944 versus time. The data are binned in 1-minute intervals. The horizontal line represents $I(>200 \text{ GeV})$ observed (Aharonian et al. 2006) from the Crab Nebula. The curve is the fit to these data of the superposition of five bursts (see text) and a constant flux.

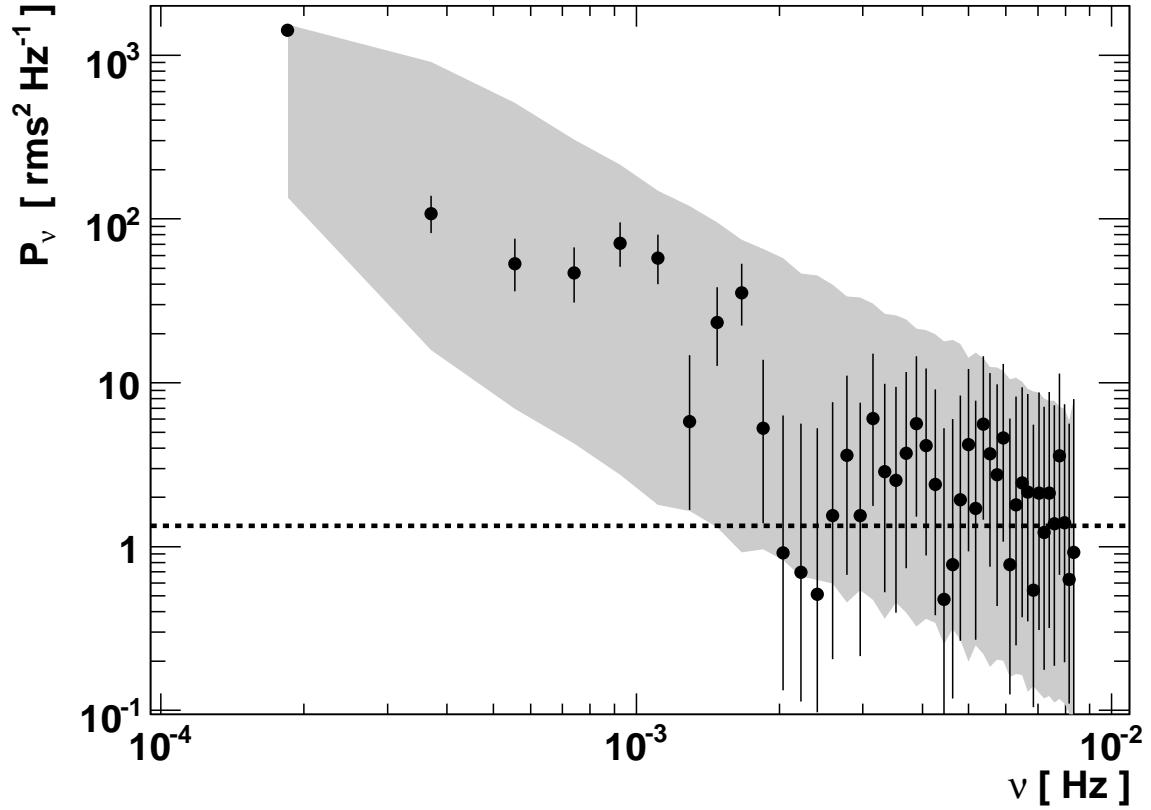


FIG. 2.— The Fourier power spectrum of the light curve and associated measurement error. The grey shaded area corresponds to the 90% confidence interval for a light curve with a power-law Fourier spectrum $P_\nu \propto \nu^{-2}$. The horizontal line is the average noise level (see text).

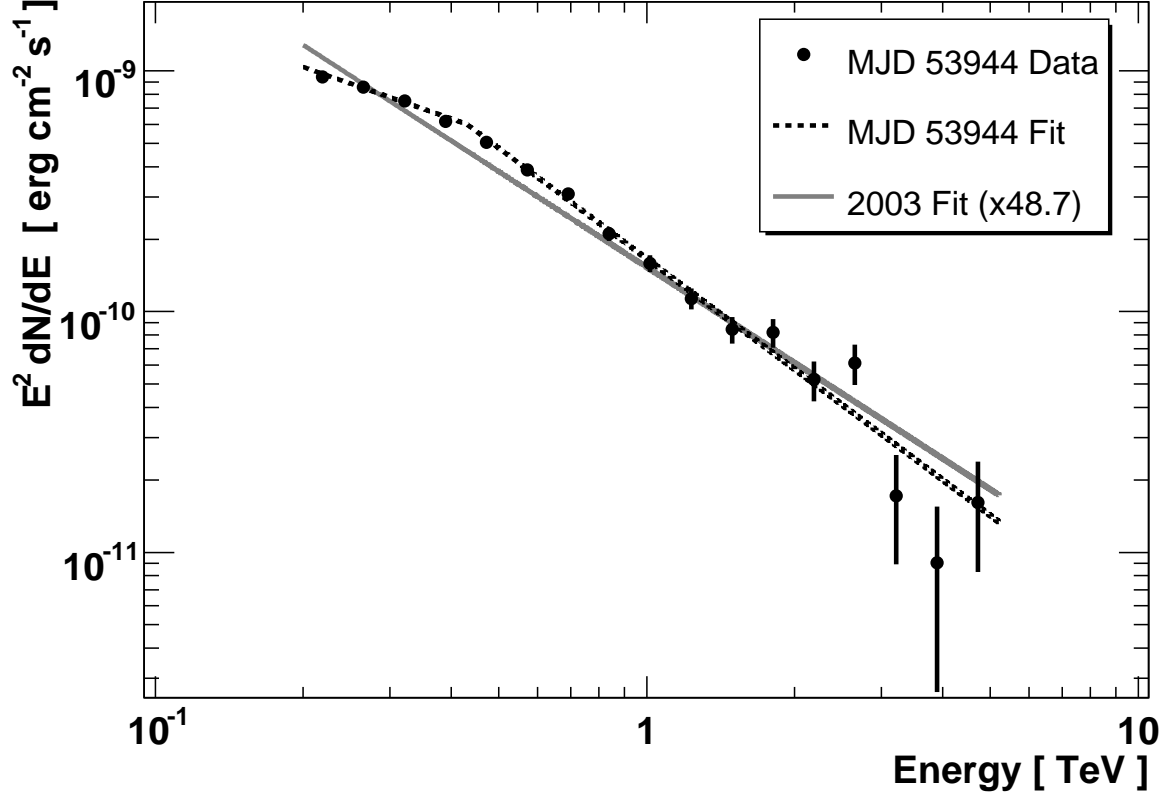


FIG. 3.— The time-averaged spectrum observed from PKS 2155–304 on MJD 53944. The dashed line is the best χ^2 fit of a broken power law to the data. The solid line represents the fit to the time-averaged spectrum of PKS 2155–304 from 2003 (Aharonian et al. 2005a) scaled by 48.7. Neither spectrum is corrected (see, e.g., Aharonian et al. 2005b) for the absorption of VHE γ -rays on the Extragalactic Background Light.